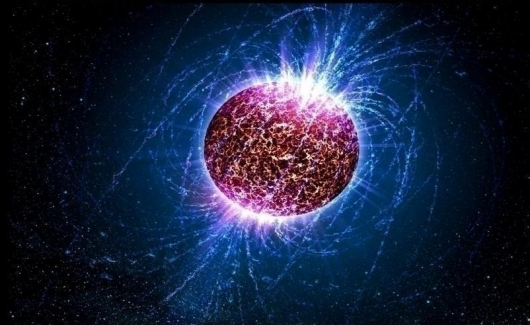


Constraints on the fermionic dark matter from observations of neutron stars

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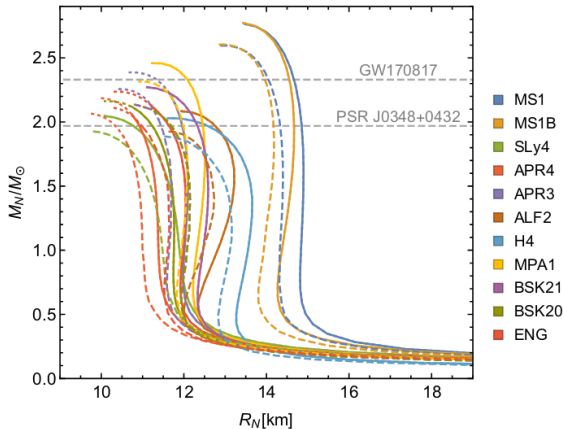
DM candidates

DM admixed
NS
Accretion onto
the NS
Conclusions



credits: Symmetry magazine

Effect of DM on NS properties



DM core contributing to 5% of the total NS mass

$$\sqrt{\sigma_D}/m_D^3 = 0.05 \text{ GeV}^{-2}$$

M. Deliyergiyev et al., PRD 99, 063015 (2019)

A. Del Popolo et al., Phys. D. Univ. 28, 100484 (2020)

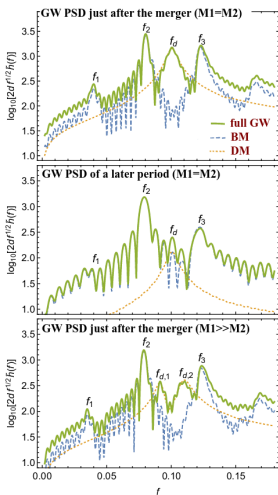
J. Ellis et al., PRD 97, 123007 (2018)

Effect of DM on GW waveform

DM admixed NS

Accretion onto the NS

Conclusions



The DM cores may produce a supplementary peak in the characteristic GW spectrum of NS mergers, which can be clearly distinguished from the features induced by the baryon component

J. Ellis et al., PLB, 781, 607 (2018)

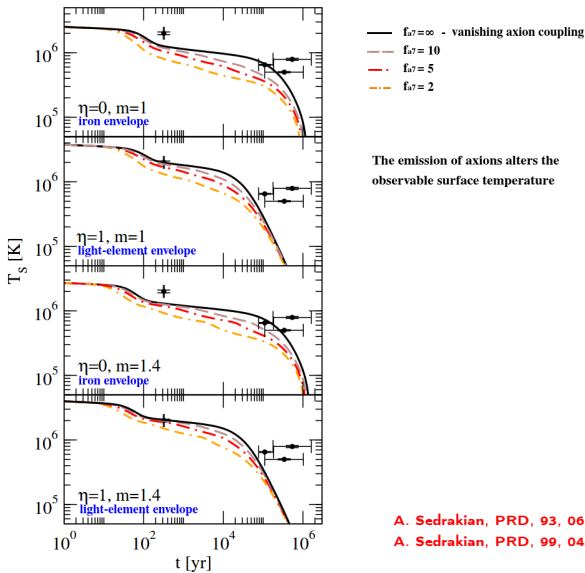
M. Bezares et al., PRD, 100, 044049 (2019)

Cooling of NS with DM

DM admixed NS

Accretion onto the NS

Conclusions



The emission of axions alters the observable surface temperature

A. Sedrakian, PRD, 93, 065044 (2016)

A. Sedrakian, PRD, 99, 043011 (2019)

2 NSs with mass above $2M_{\odot}$

- PSR J0348-0432: $M = 2.01^{+0.04}_{-0.04} M_{\odot}$ (Antoniadis et al.'13)
- PSR J0740+6620: $M = 2.14^{+0.20}_{-0.18} M_{\odot}$ (Cromartie et al.'19)

Dark matter EoS

- **Asymmetric dark matter**
relativistic Fermi gas of noninteracting particles with the spin 1/2
A. Nelson, S. Reddy, D. Zhou, JCAP 07, 012 (2019)

Baryon matter EoS

- **EoS with induced surface tension (IST EoS)**
consistent with:
nuclear matter ground state properties,
proton flow data,
heavy-ion collisions data,
astrophysical observations,
tidal deformability constraint from the NS-NS merger (GW170817)
VS, I. Lopes, A. Ivanytskyi, ApJ, 871, 157 (2019)
VS, A. Ivanytskyi, K. Bugaev, et al., NPA, 924, 24 (2014)

TOV equations

2 TOV equations:

$$\frac{dp_B}{dr} = -\frac{(\epsilon_B + p_B)(M + 4\pi r^3 p)}{r^2(1 - 2M/r)}$$

$$\frac{dp_D}{dr} = -\frac{(\epsilon_D + p_D)(M + 4\pi r^3 p)}{r^2(1 - 2M/r)}$$

BM and DM are coupled only through gravity, and their energy-momentum tensors are conserved separately

total pressure $p(r) = p_B(r) + p_D(r)$

gravitational mass $M(r) = M_B(r) + M_D(r)$, where $M_j(r) = 4\pi \int_0^r \epsilon_j(r') r'^2 dr'$ (j=B,D)

Fraction of DM inside the star:

$$f_x = \frac{M_D(R_D)}{M_T}$$

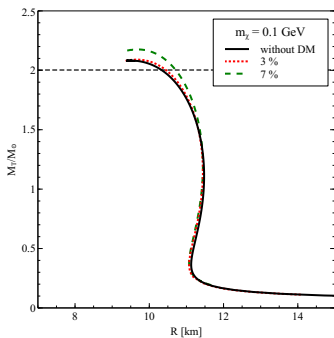
$M_T = M_B(R_B) + M_D(R_D)$ - total gravitational mass

Mass-Radius diagram of the DM admixed NSs

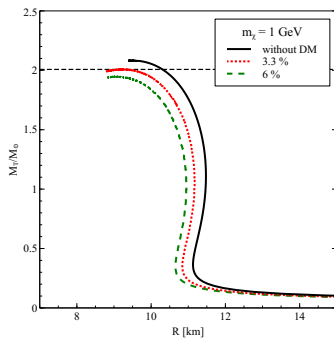
DM admixed NS

Accretion onto the NS

Conclusions



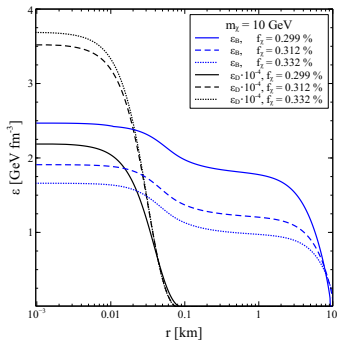
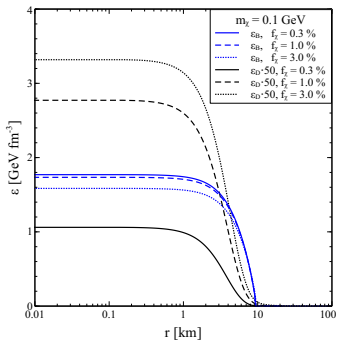
$M_{max} > 2 M_{\odot}$ for any f_{χ}



for $f_{\chi} = 3.3\%$ M_{max} equals to $2 M_{\odot}$
 further increase of the DM fraction
 leads to $M_{max} < 2 M_{\odot}$

O. Ivanytskyi, VS, I. Lopes, PRD, 102, 063028 (2020)

Internal structure of the stars



$R_D = 9.4 \text{ km}$ for $f_\chi = 0.3\%$
 $R_D = 21.2 \text{ km}$ for $f_\chi = 1.0\%$
 $R_D = 135.2 \text{ km}$ for $f_\chi = 3.0\%$

Large values of R_D relate to the existence of dilute and extended halos of DM around a baryon core of NS

DM admixed NS

Accretion onto the NS

Conclusions

Dark matter and neutron star structure

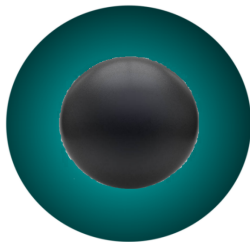
DM admixed
NS

Accretion onto
the NS

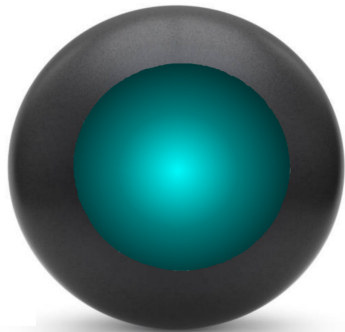
Conclusions



dark matter core



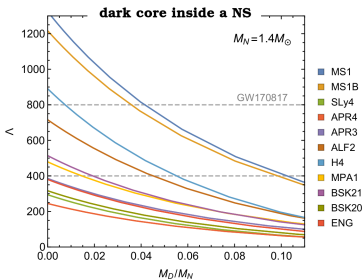
dark core inside a NS



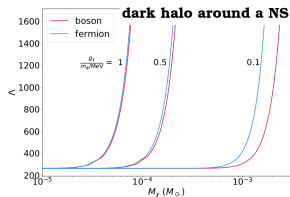
dark halo around a NS

Tidal deformability can be used as a probe of non standard physics!

- DM modifies tidal interactions between merging NSs in an opposite way: increases due to the presence of DM halo and decreases in the case of a DM core



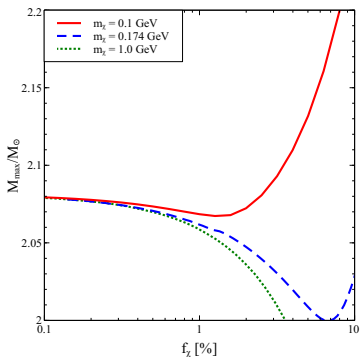
J. Ellis et al., PLB, 781, 607 (2018)



λ increases rapidly with increasing total DM mass M_χ . For self-interacting DM with $g_\chi/m_\phi > 1 \text{ MeV}^{-1}$, $M_\chi > 10^{-4} M_\odot$ will boost λ above the upper bound (≈ 800) set by GW170817.

A. Nelson et al., JCAP 7, 12 (2019)

Maximal mass of NS as a function of the DM fraction



for $m_\chi = 0.174$ GeV M_{max} is $2 M_\odot$

DM particles with $m_\chi \leq 0.174$ GeV are consistent with the $2 M_\odot$ constraint for any f_χ

For heavier DM particles the NS mass can reach $2 M_\odot$ only if f_χ is limited from above

DM admixed
NS

Accretion onto
the NS

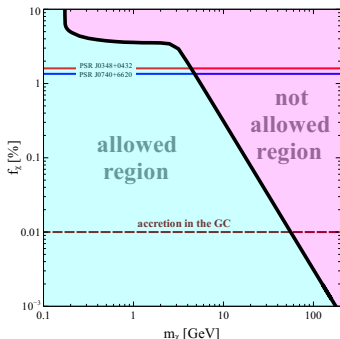
Conclusions

Constraint on the mass of DM particles

DM admixed
NS

Accretion onto
the NS

Conclusions



Navarro-Frenk-White distribution for DM:

$$\rho_{\chi}(d) = \rho_c \cdot \frac{d_c}{d} \cdot \left(1 + \frac{d}{d_c}\right)^{-2} \quad (1)$$

$$\rho_c = 5.22 \pm 0.46 10^7 M_{\odot} \text{kpc}^{-3} \text{ and } d_c = 8.1 \pm 0.7 \text{ kpc}$$

H.-N. Lin, X. Li, arXiv:1906.08419 (2019)

BM distribution in a stellar disc:

$$\rho_B(d) = \rho_{dc} e^{-\frac{d}{d_{dc}}} \quad (2)$$

$$\rho_{dc} = 15.0 M_{\odot} \text{pc}^{-3} \text{ and } d_{dc} = 3.0 \text{ kpc}$$

Y. Sofue, Publ. Astr. Soc. Jap., 65, 118 (2013)

| Pulsar | distance to the GC | f_{χ}^* (NFW distr) | f_{χ}^* (Einasto distr) |
|----------------|--------------------|--------------------------|------------------------------|
| PSR J0348+0432 | 9.9 kpc | $1.6 \pm 0.4 \%$ | $1.35 \pm 0.05 \%$ |
| PSR J0740+6620 | 8.6 kpc | $1.35 \pm 0.35 \%$ | $1.12 \pm 0.049 \%$ |

f_{χ}^* corresponds the DM fraction in the surrounding medium around the NS

Fraction inside the NS will depend on the accretion rate during all the life stages of a star and the cross-section of DM with BM

DM accumulation regimes

- **Progenitor**

During the star formation stage the initial mixture of DM and BM contracting to form the progenitor star. Trapped DM undergoes scattering processes with baryons leading to its kinetic energy lost and thermalisation.

- **Main sequence (MS) star**

From this stage of star evolution accretion rate increases due to big gravitational potential of the star. In the most central Galaxy region $M_{acc} \approx 10^{-5} M_{\odot} - 10^{-9} M_{\odot}$.

- **Supernova explosion & formation of a proto-NS**

The newly-born NS will be surrounded by the dense cloud of DM particles with the temperature and radius that corresponds to the last stage of MS star evolution, i.e. a star with a silicone core.

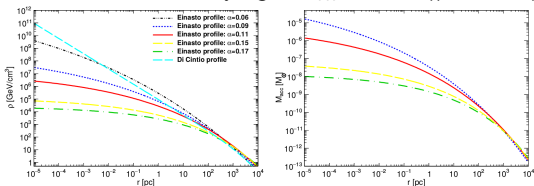
Kouvaris & Tinyakov (2010)

In addition, a significant amount of DM can be produced during the supernova explosion and mostly remain trapped inside the star.

- **Equilibrated NS**

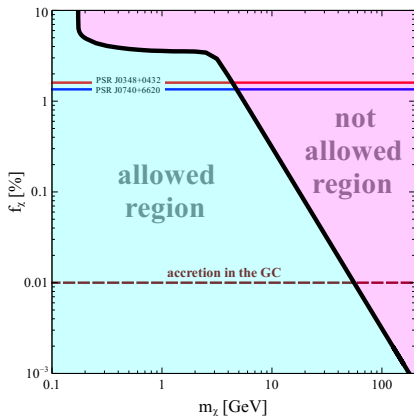
$$M_{acc} \approx 10^{-14} \left(\frac{\rho_{\chi}}{0.3 \frac{\text{GeV}}{\text{cm}^3}} \right) \left(\frac{\sigma_{\chi n}}{10^{-45} \text{cm}^2} \right) \left(\frac{t}{\text{Gyr}} \right) M_{\odot}, \quad (3)$$

In the most central Galaxy region $M_{acc} \approx 10^{-5} M_{\odot} - 10^{-8} M_{\odot}$.



A. Del Popolo, et al. (2019)

DM constraint in the GC



$2 M_{\odot}$ NS in the GC $\Rightarrow m_{\chi} < 60$ GeV

More precise modeling of DM accumulation inside the NSs will put more tight constraints on m_{χ} .

- We demonstrated that DM lighter than 0.2 GeV can create an extended halo around the NS leading not to decrease but to increase of the NS total (gravitational) mass.



Credit: NASA Goddard

- By using recent results on the distribution of DM in Milky Way, we made an estimation of the fraction of DM in NSs in the GC. Measurements of a $2 M_{\odot}$ NS in the GC will impose an upper constraint on the mass of DM particles of ~ 60 GeV.
- We expect to have more NSs observations and measurements of their masses with higher precision from the following telescopes:

radio telescopes

- the Karoo Array Telescope (MeerKAT)
- the Square Kilometer Array (SKA)
- the Next Generation Very Large Array (ngVLA)

space telescopes

- the Neutron Star Interior Composition Explorer Mission (NICER)
- the Advanced Telescope for High Energy Astrophysics (ATHENA)
- the enhanced X-ray Timing and Polarimetry mission (eXTP)
- the Spectroscopic Time-Resolving Observatory for Broadband Energy X-rays (STROBE-X)

Thanks for your attention!

DM admixed
NS

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the NS

Conclusions

