LATTES Science Case

The 40 – 500 GeV energy domain

Giovanni La Mura

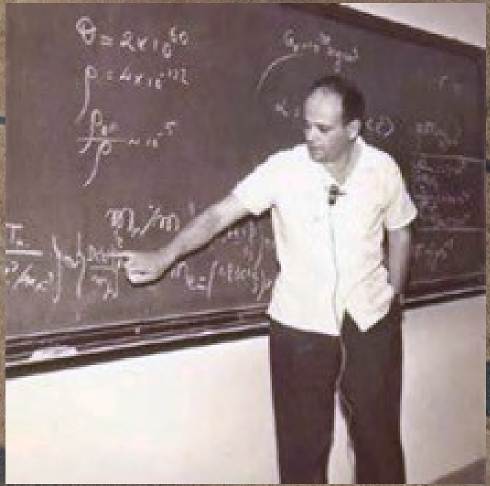
on behalf of the Fermi-LAT Collaboration
and of the LATTES team

LIP

7th LATTES Meeting, Coimbra, September 28th 2018

Gamma-ray Astrophysics in multi-messenger era

- **Status of gamma-ray observations**
- Galactic Astrophysics
- Extragalactic Astrophysics
- How LATTES can contribute
- Concluding remarks



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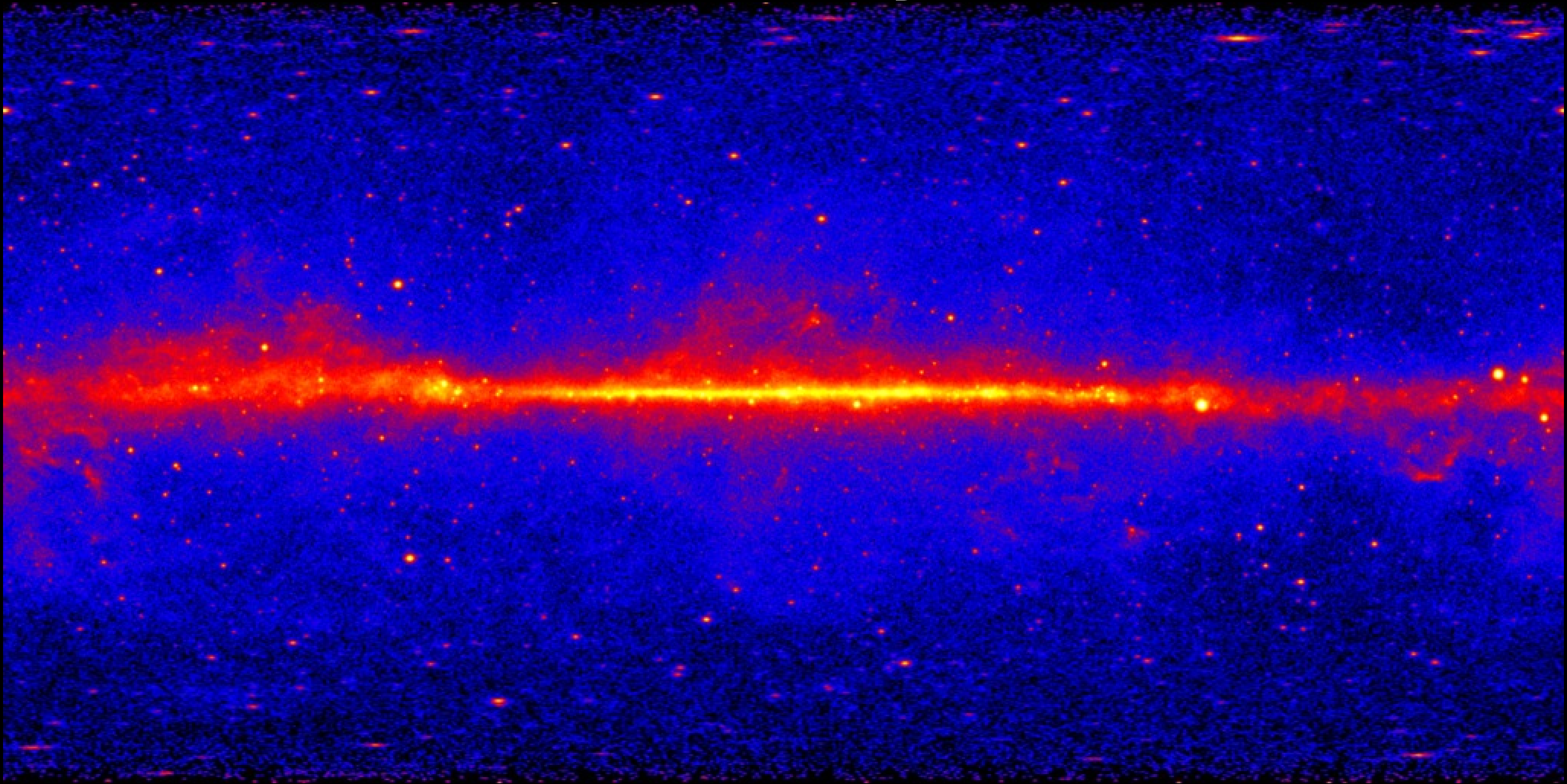
Status of Gamma-ray Observations



All-sky Cartesian projection of the Fermi-LAT gamma-ray observations above 1 GeV in photon energy, achieved after 5 years of monitoring. This is, to date, the **deepest published view** of the High-Energy sky.
Image credit: NASA/DOE/Fermi LAT Collaboration

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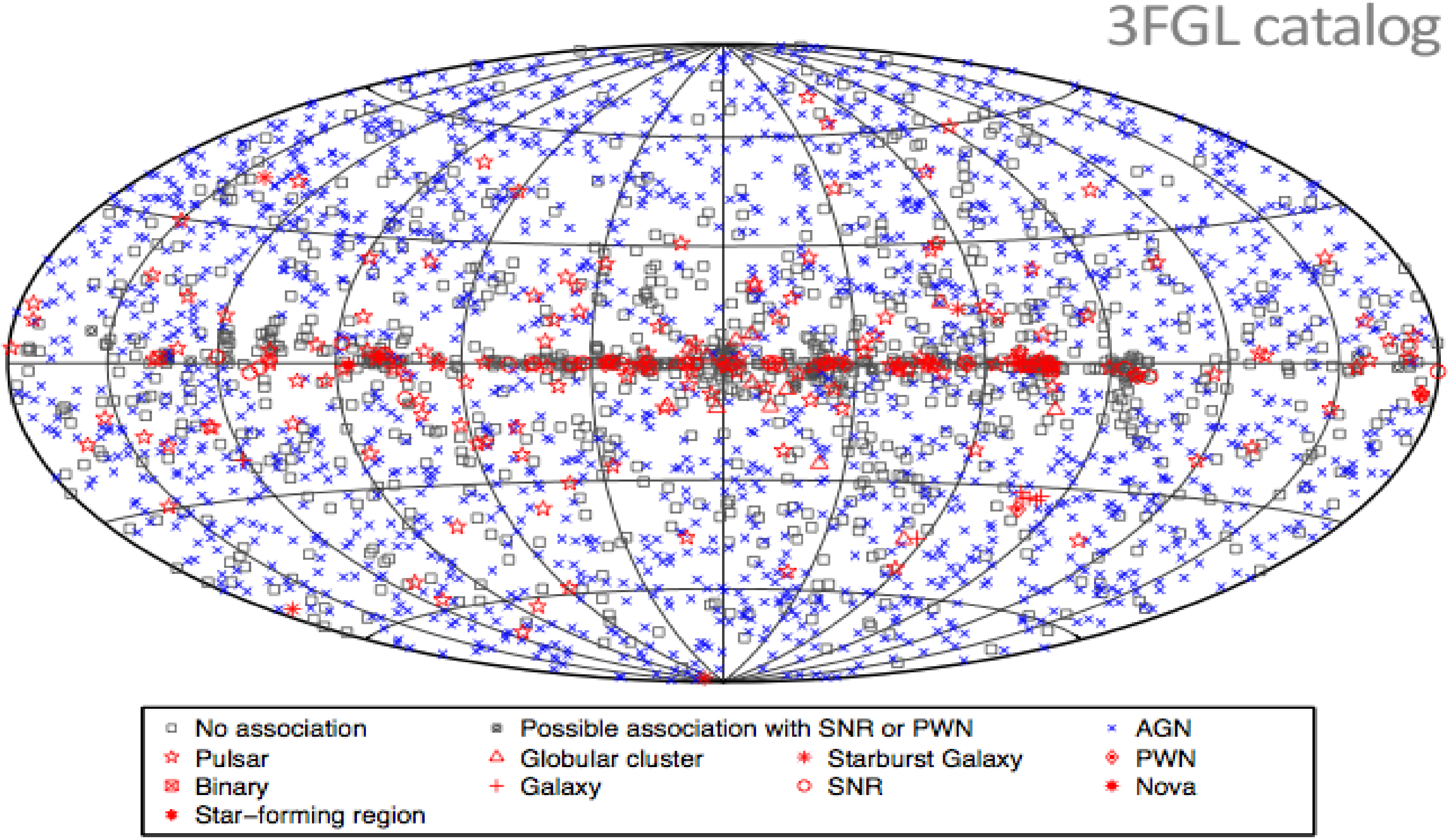

Distribution of the 3033 Fermi/LAT sources in the 3FGL Catalog (Acero et al. 2015, ApJS, 218, 23), with the corresponding classification. The sources are detected with 100 MeV < E < 300 GeV. Image from ASDC.

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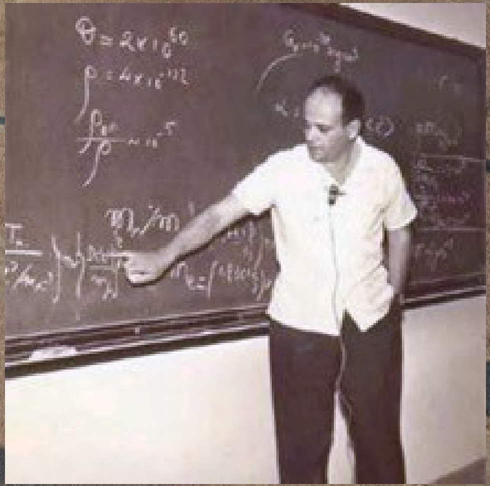

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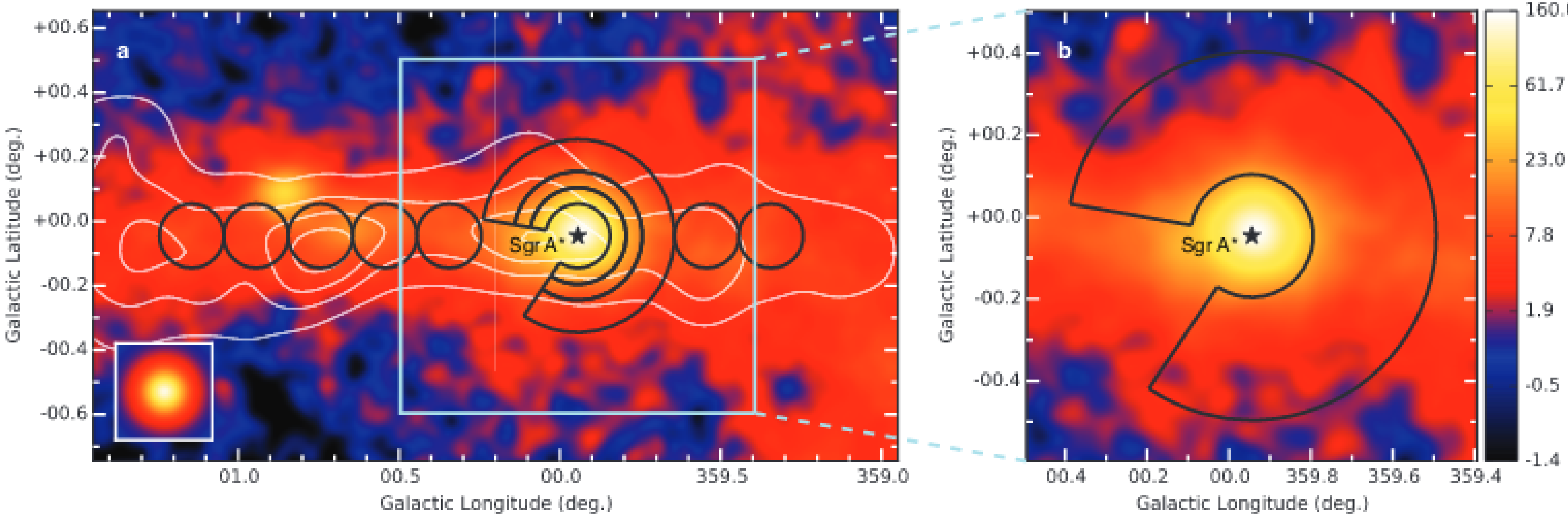
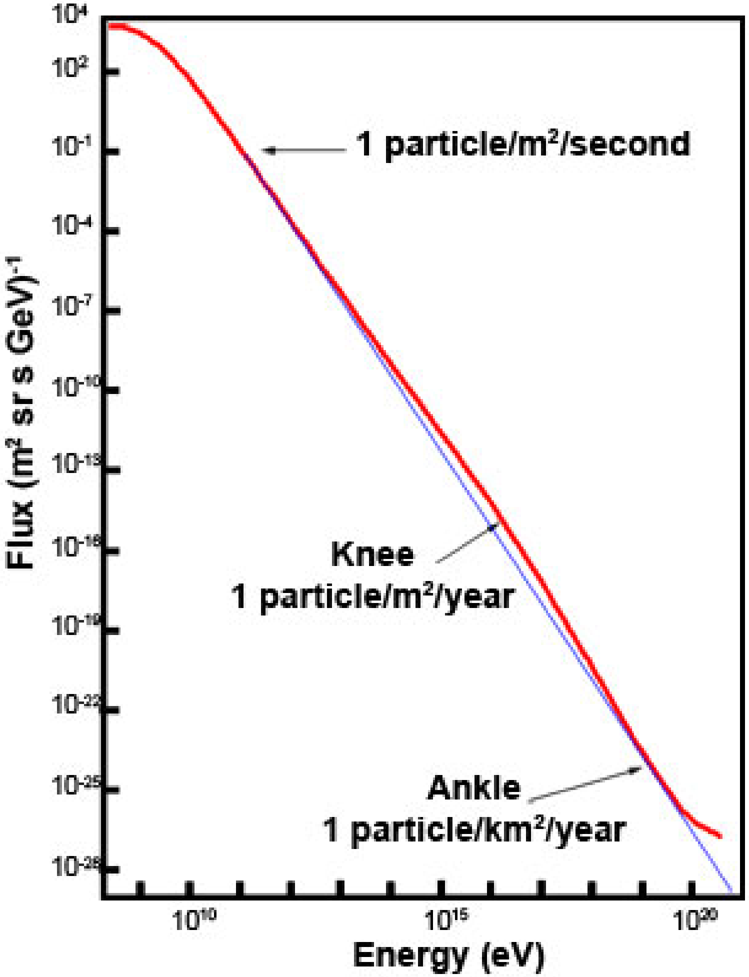
The Galactic Centre

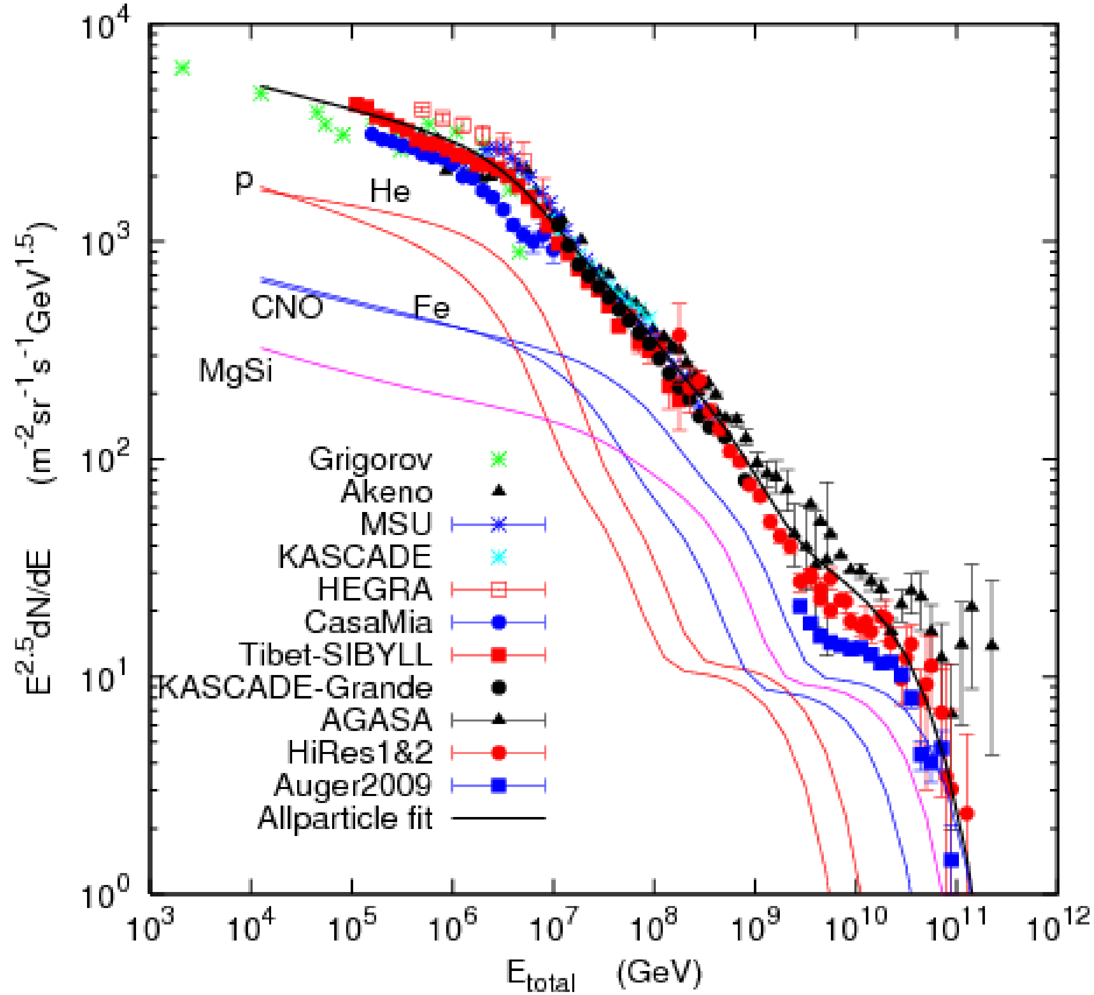


A view of the Galactic Centre in the TeV domain, produced by a scan performed with HESS (HESS Coll. 2016, Nature, 531, 476).

The Galactic Centre is an extremely complex region, expected site of the highest Dark Matter density in our vicinity, but also populated by stellar evolution remnants (neutron stars), the Sgr A* black hole, and dense clouds of gas, interacting with cosmic rays. It is identified as the site of a Pevatron (accelerator of particles up to the PeV domain).

Acceleration of Cosmic Rays





A full understanding of the acceleration sites of cosmic rays and their connection with the production of gamma rays is also related to the observed spectral features and to the possibility to extensively monitor the Galactic Plane and the Galactic Centre.

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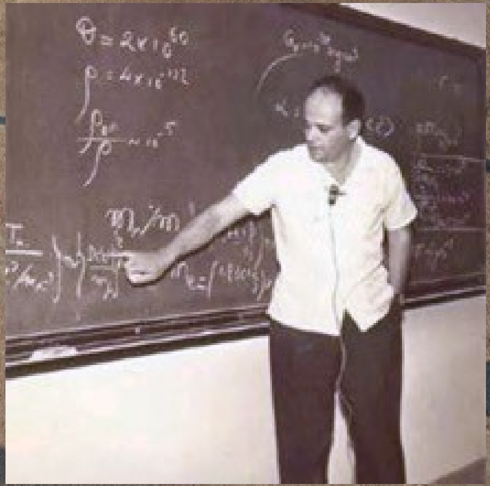

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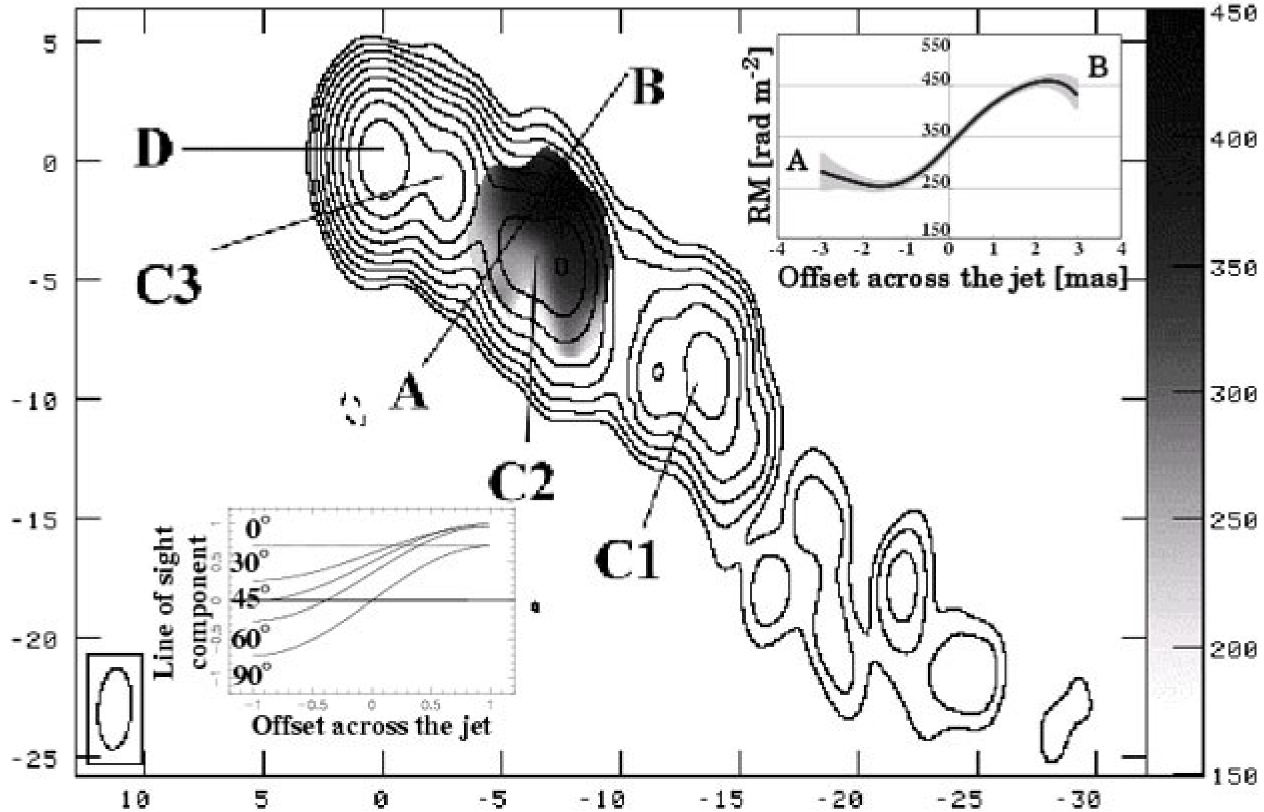
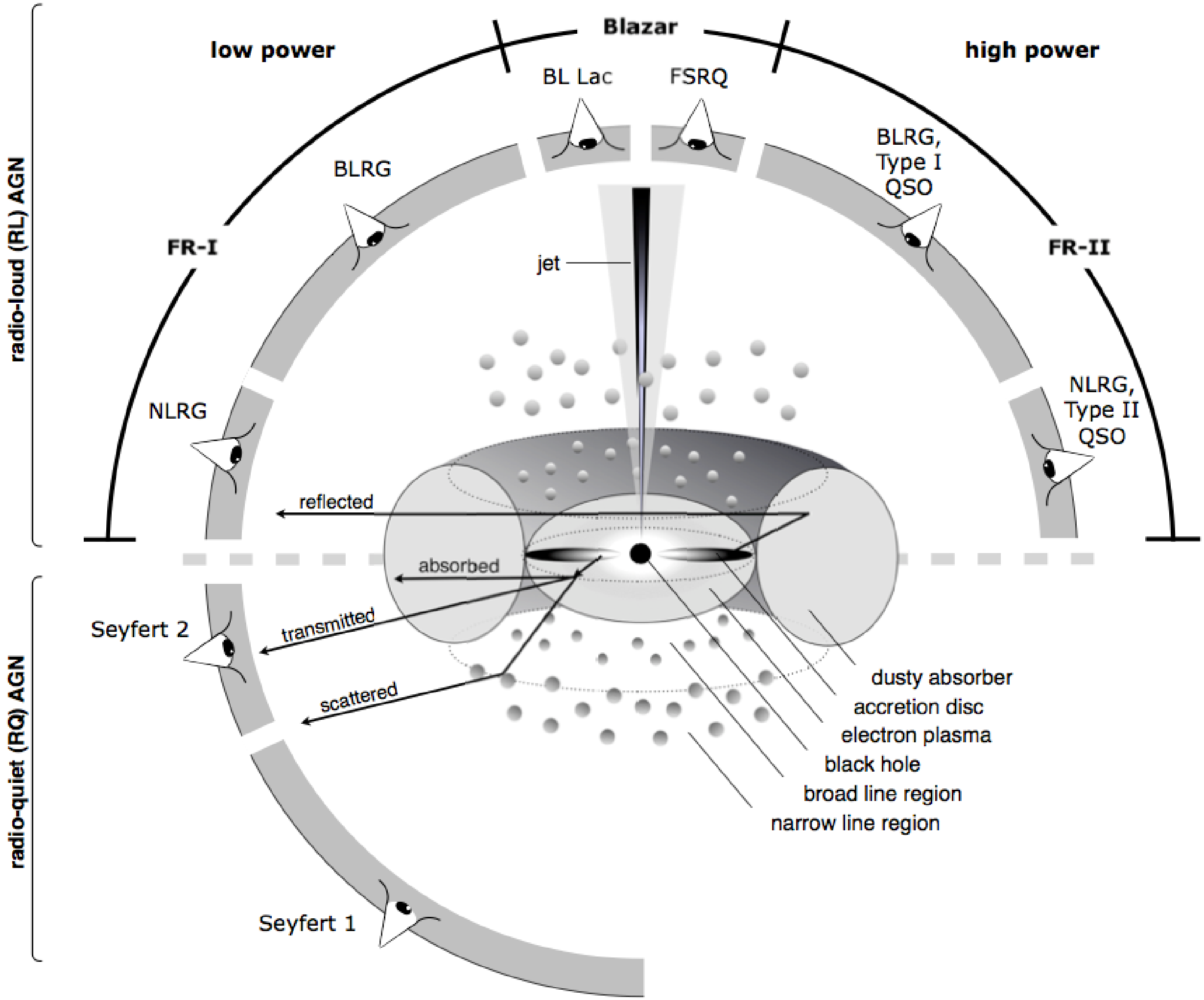
Jets from Active Galactic Nuclei



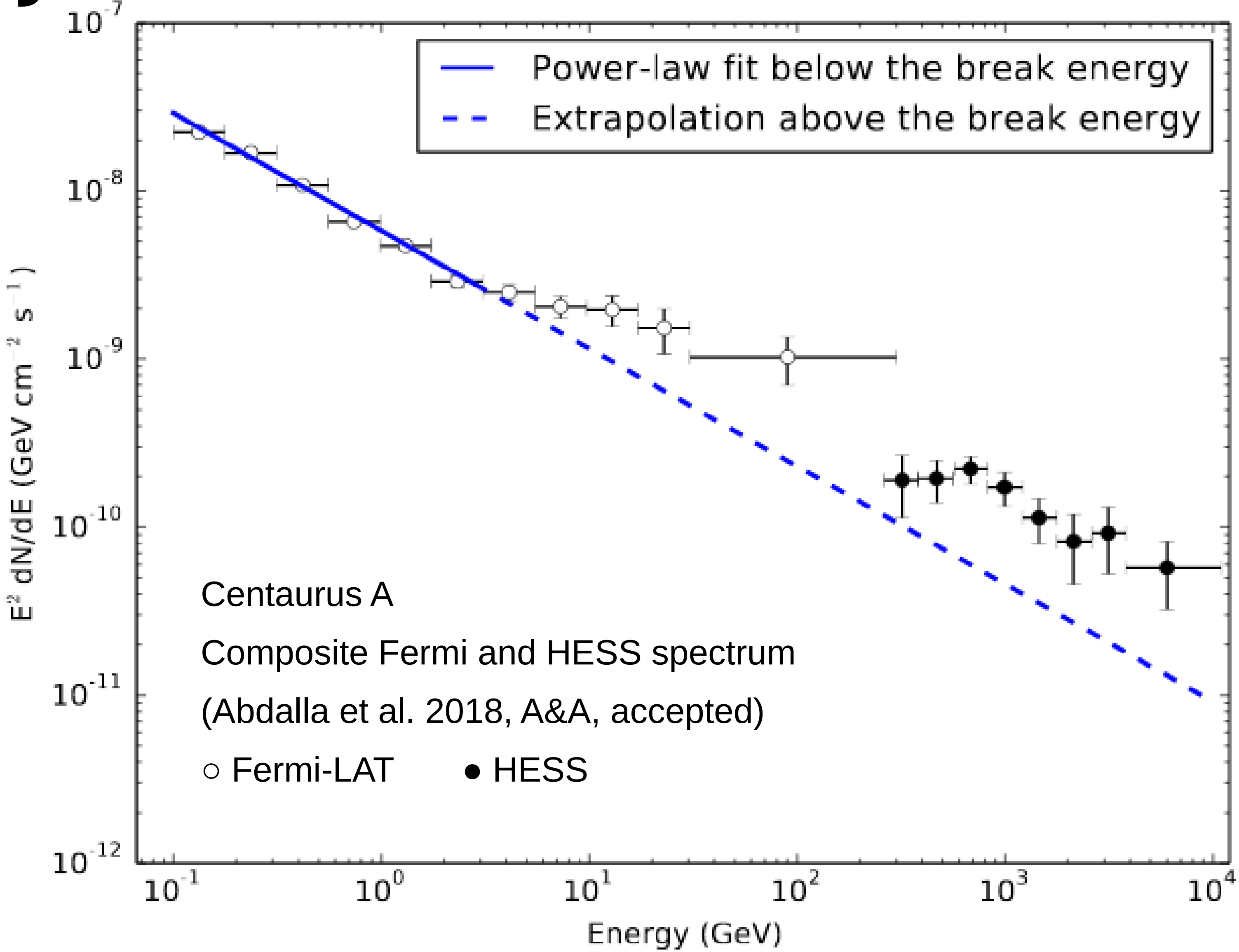
The **helical structure** of the magnetic field *B* in the jet of 3C 273, as detected by the variation of the **Faraday Rotation Measure** (RM, grey scale), super-imposed to the jet radio intensity map. The RM profile is extracted along the line of sight and across the jet, showing a twisted structure that takes part to the confinement of the jet plasma.

Jets from Active Galactic Nuclei



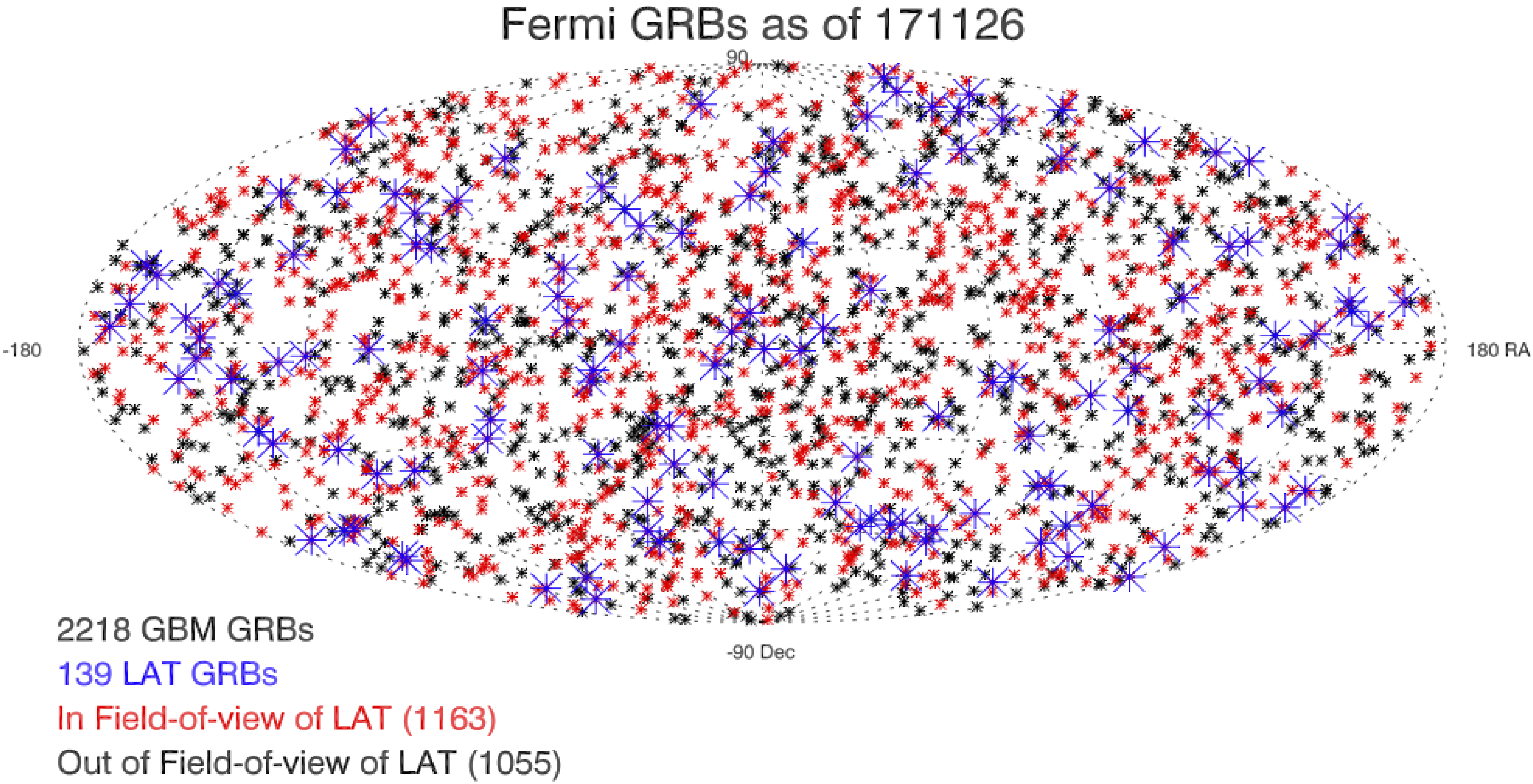
The relativistic jet is directly observed in hundreds of AGNs. It is the production site of most of the radio, X-ray and gamma-ray emission. It is a prominent feature of QUASARs and Radio Galaxies and, when seen along the axis, results in the form of a BLAZAR.

Jets from Active Galactic Nuclei



Jet Physics is not a proprietary field of BLAZARs: misaligned sources, such as Radio Galaxies, are fundamental to understand how the jet radiates. A monitoring instrument is necessary for this.

Gamma Ray Bursts



Gamma-ray Bursts are flashes of radiation occurring randomly in the sky. They last from less than 1s up to some hours. There is, on average, more than 1 event per day and their energy peak is typically in the sub-MeV domain.

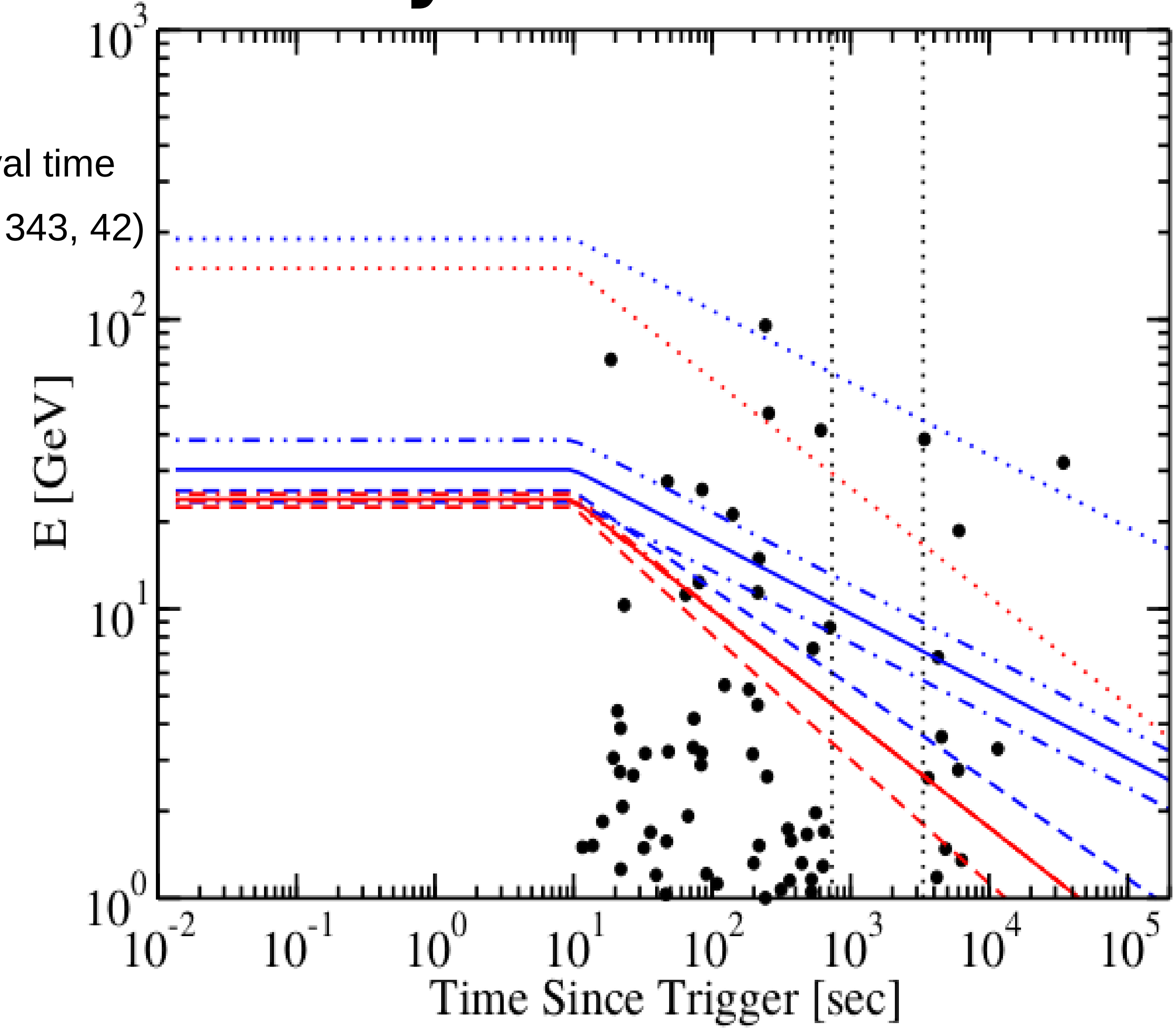
Credit: Fermi-LAT and GBM Collaborations

Gamma Ray Bursts

GRB 130427A

Photon energy as function of arrival time

(Ackermann et al. 2013, Science, 343, 42)



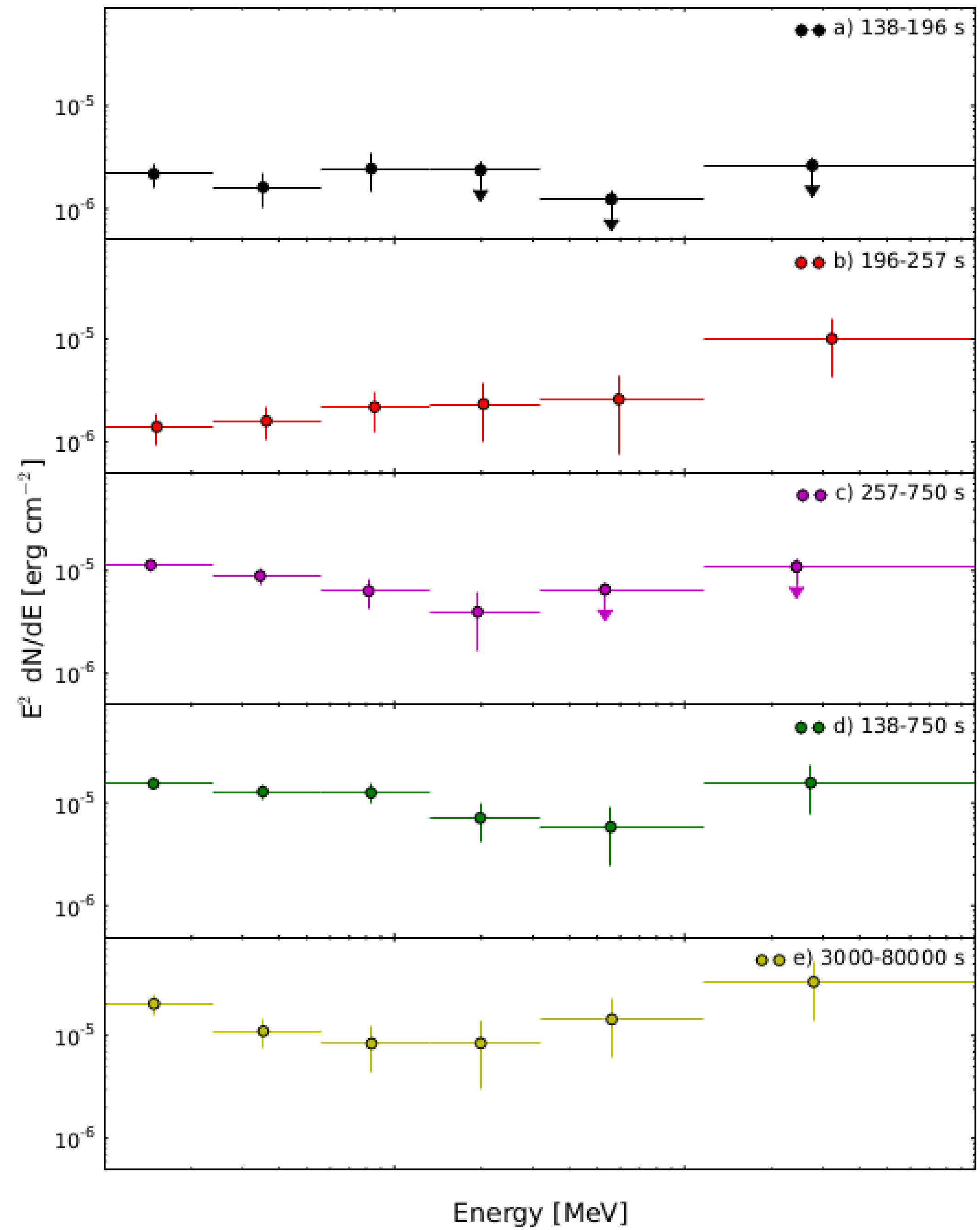
Some Gamma-ray Bursts have been shown to produce a glow of high-energy emission, typically tens to hundreds seconds after the main event trigger. This is likely due to the GRB blast wave interactions.

Gamma Ray Bursts

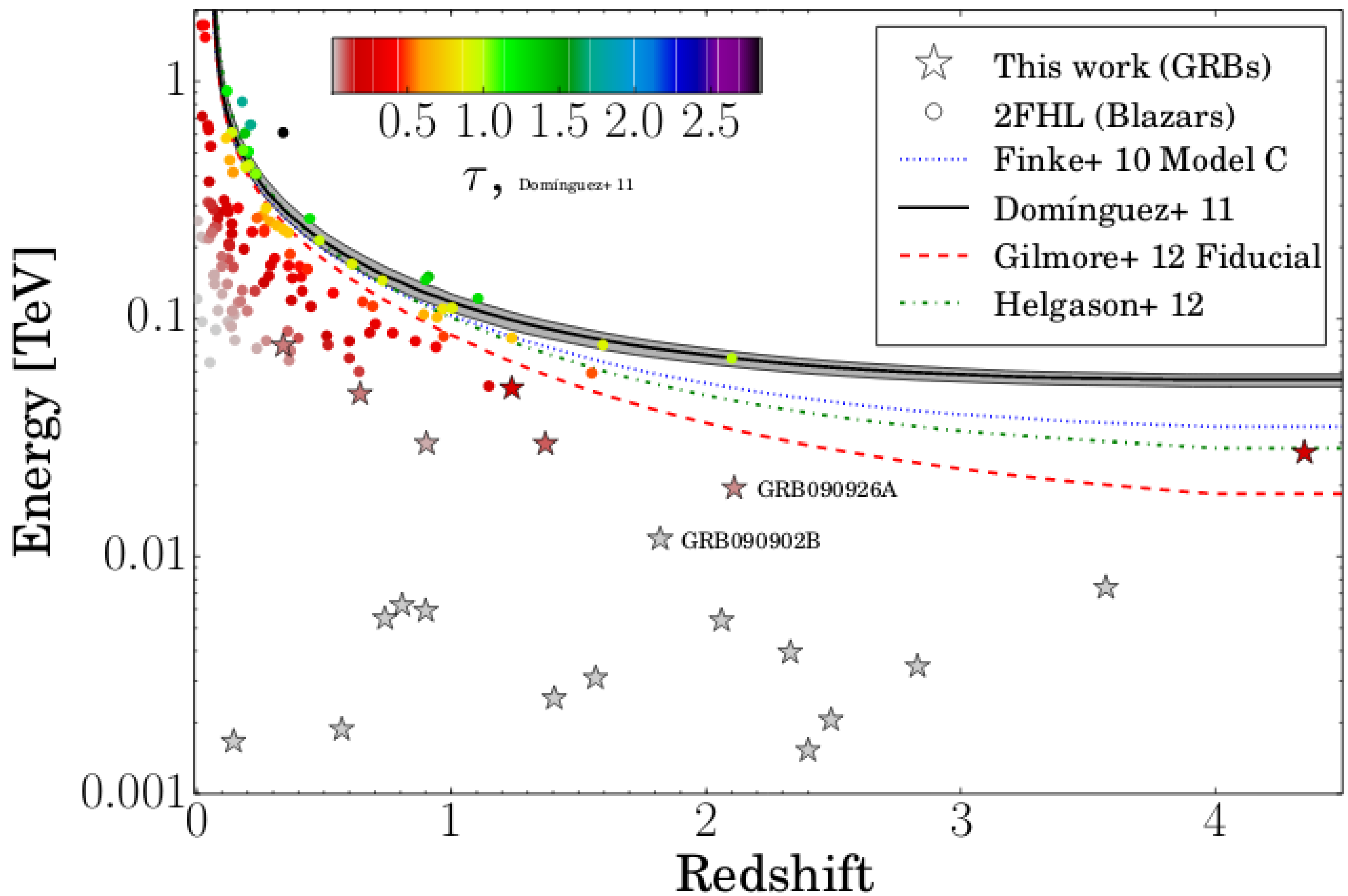
SED of GRB 130427A in different time bins.

(Ackermann et al. 2013, Science, 343, 42)

We are currently lacking a large FoV instrument that may assess the existence of spectral breaks, which, for events of known redshift, are directly connected with the Universe's gamma-ray opacity.



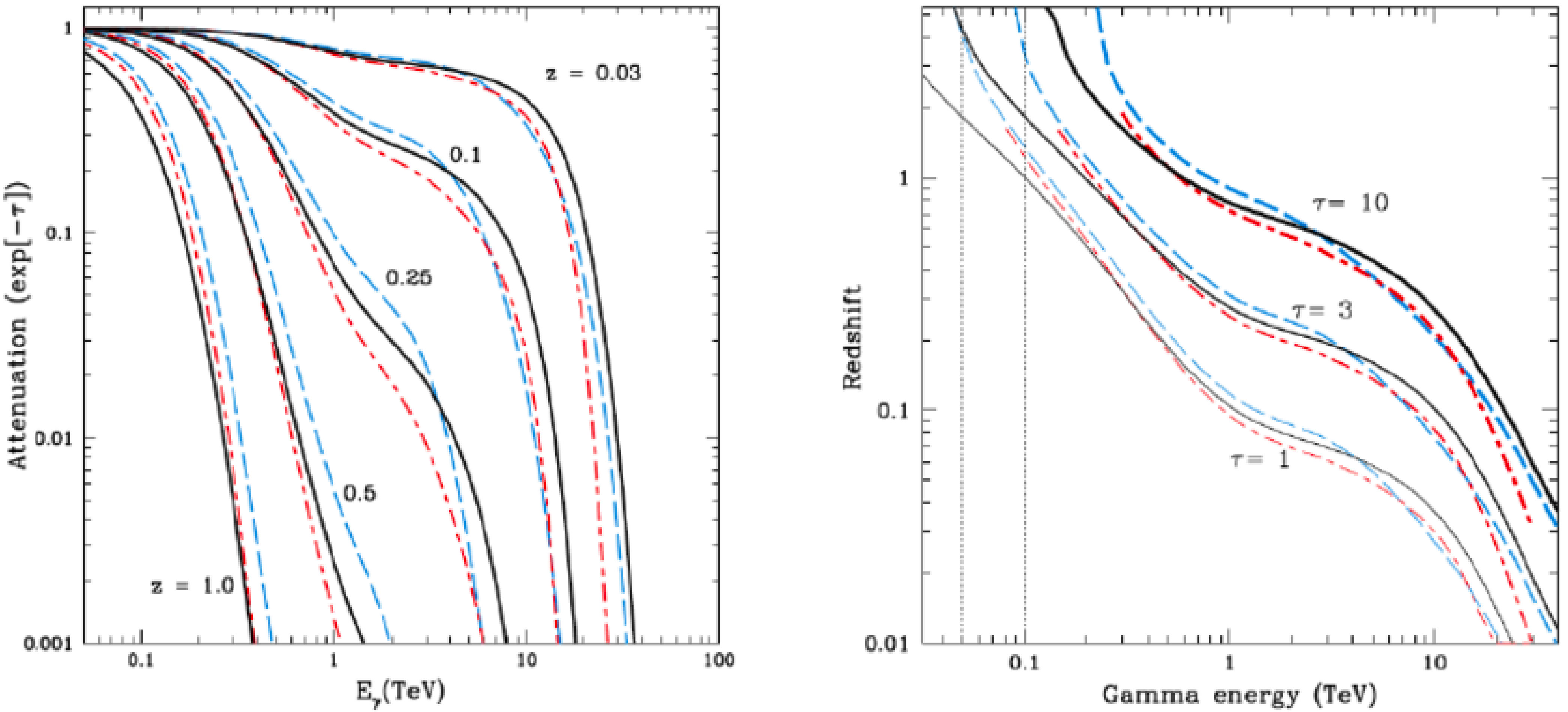
Gamma Ray Bursts



Maximum Photon Energy vs. Redshift distribution for blazars and GRBs detected by Fermi-LAT. The data show that GRBs can actually produce photons with E > 10 GeV.

Figure from Desai et al. 2017, ApJ, 850, 73

Gamma Ray Bursts



The sub-TeV energy domain is crucial to discriminate among the different models of Extragalactic Background Light that predict the Universe gamma-ray opacity. Such models are a fundamental cosmological constraint.

(Fig. from Primack et al. 2011; DOI: 10.1063/1.3635825)

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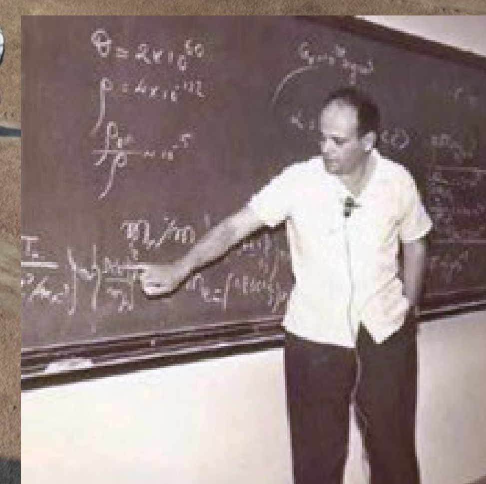

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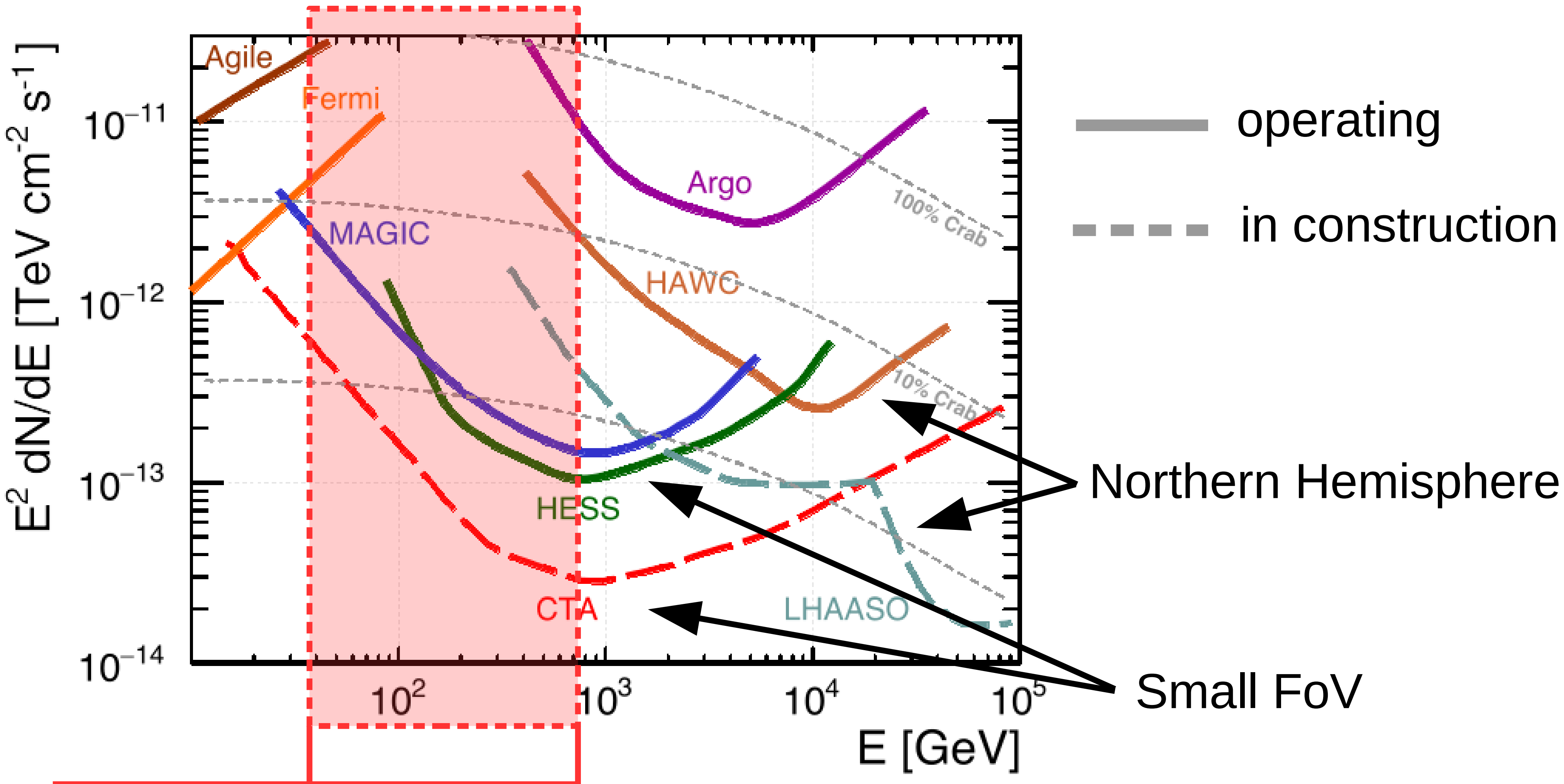
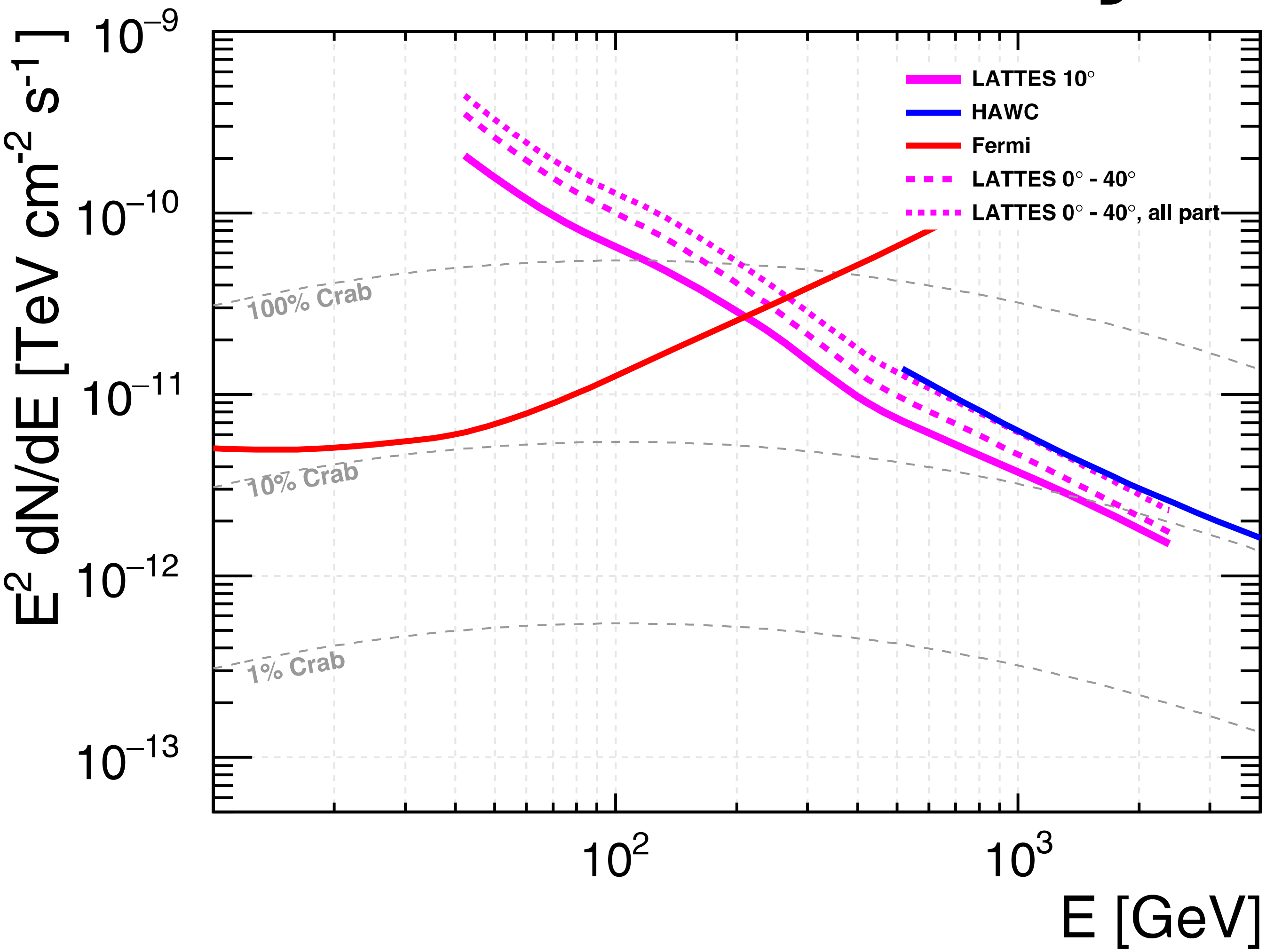
Monitoring the spectrum is a priority!



- No wide FoV instrument working in the Southern Hemisphere
- Gap between satellite and ground-based observations

LATTES sensitivity



- Sensitivity to steady sources in one year of operation
- *The ability to reach the lowest energies is critical to most of the science discussed so far*

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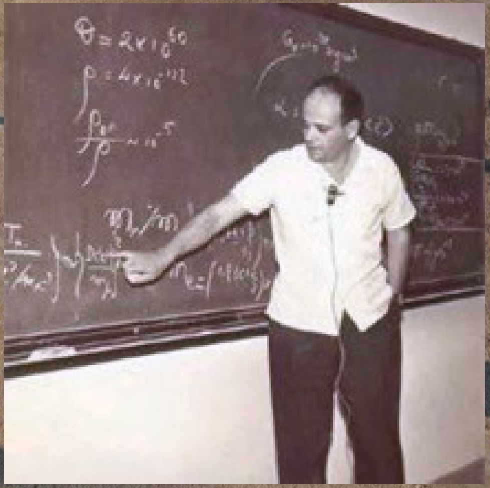

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Concluding remarks

- There is a wide population of Galactic and Extragalactic sources that should be monitored and investigated in high-energy gammas (E > 100 GeV)
- *LATTES* is ideal to monitor the Galactic Plane and Centre
- *LATTES* has very good possibilities to address questions as cosmic ray acceleration, AGN jet physics and Universe gamma-ray opacity
- Many of the scientific goals above are best pursued by including the sub-TeV domain in the instrument capability

Acknowledgements

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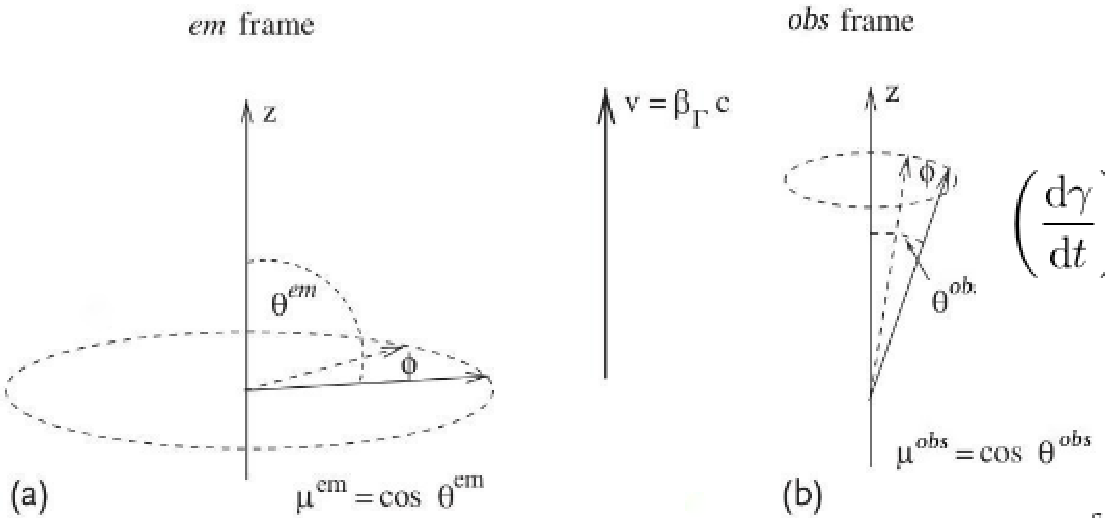
TÉCNICO LISBOA

Gamma-ray Large Area Space Telescope
GLAST

Backup slides

Physics of AGN jets

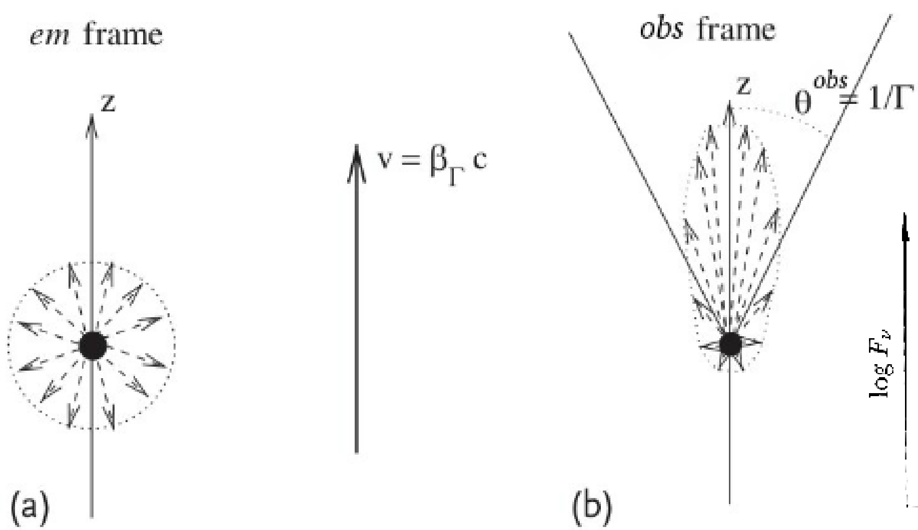


Particle energy loss

$$\left(\frac{\mathrm{d}\gamma}{\mathrm{d}t}\right)_{sy} = -\frac{4}{3}c\sigma_T\frac{u_B}{m_e c^2}Z^4\left(\frac{m_e}{m}\right)^3\beta^2\gamma^2$$

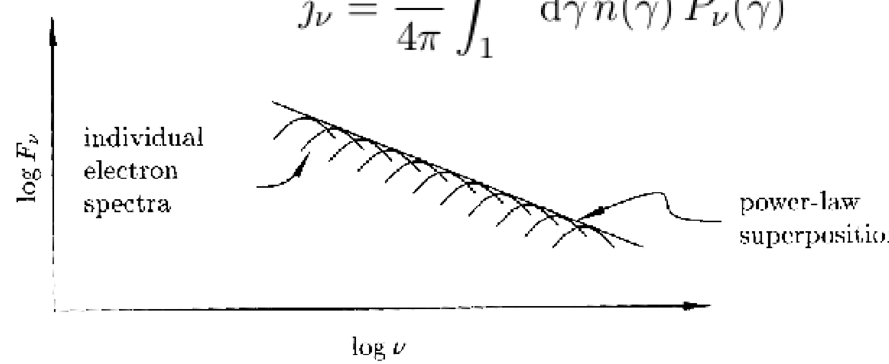
δ - approximation

$$P_\nu^{sy,\delta}(\gamma) = \frac{32\pi}{9}\left(\frac{q^2}{mc^2}\right)^2 u_B\beta^2\gamma^2\delta(\nu-\nu_c)$$



Composite spectrum

$$j_\nu = \frac{1}{4\pi}\int_1^\infty \mathrm{d}\gamma\, n(\gamma)\, P_\nu(\gamma)$$



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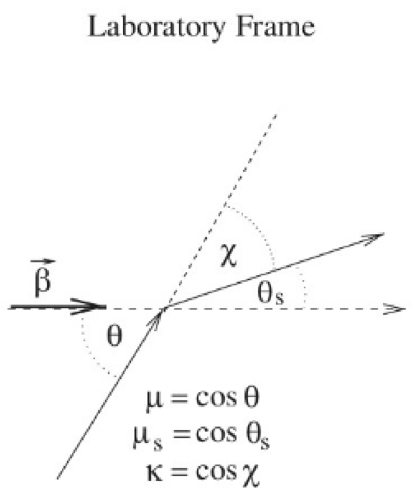
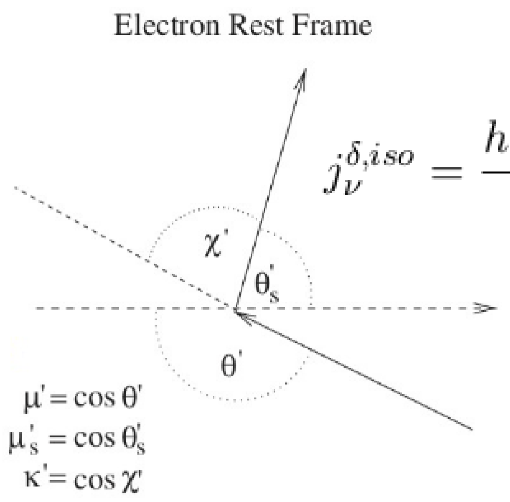
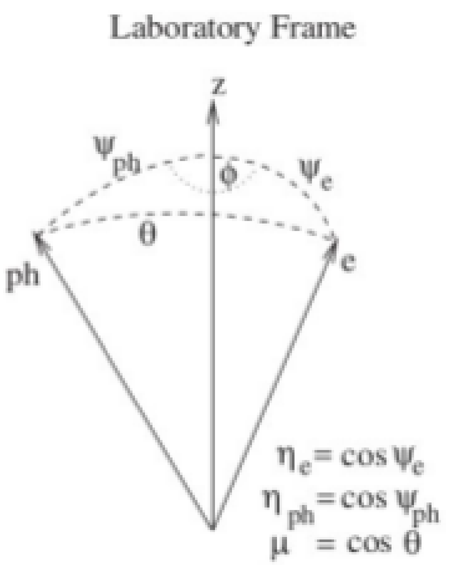


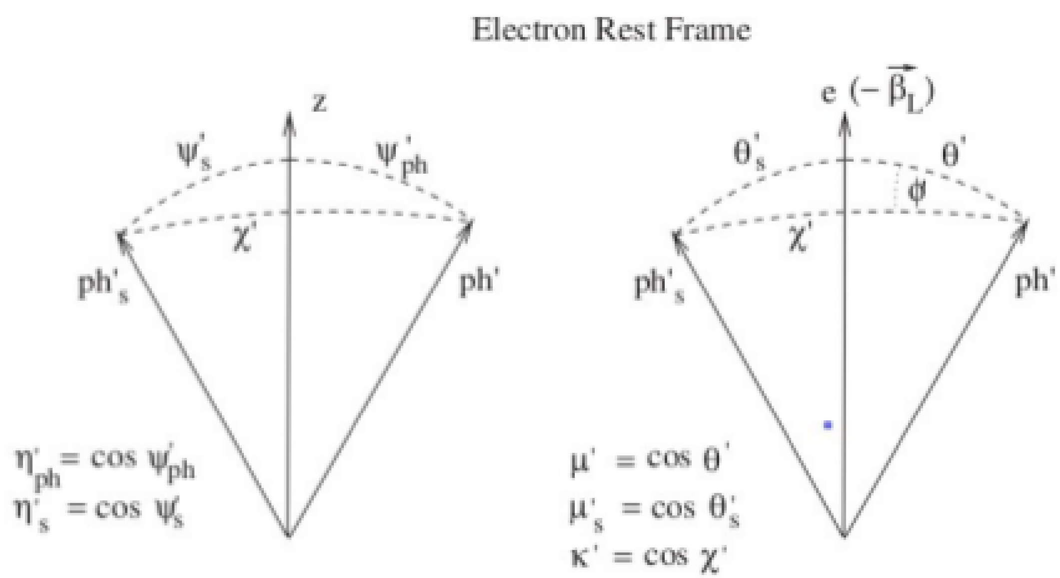


Compton emissivity

$$j_\nu^{\delta,iso} = \frac{h\, c\, \sigma_T\, \varepsilon_S^2}{8\pi} \int_{\varepsilon_S}^{\infty} \mathrm{d}\gamma\, \frac{n_e(\gamma)}{\gamma^4} \int_{\varepsilon_S/(2\gamma^2)}^{\infty} \mathrm{d}\varepsilon\, \frac{n_{ph}(\varepsilon)}{\varepsilon^2}$$

$$L_{IC} / L_{sync} \approx u_{ph} / u_B$$





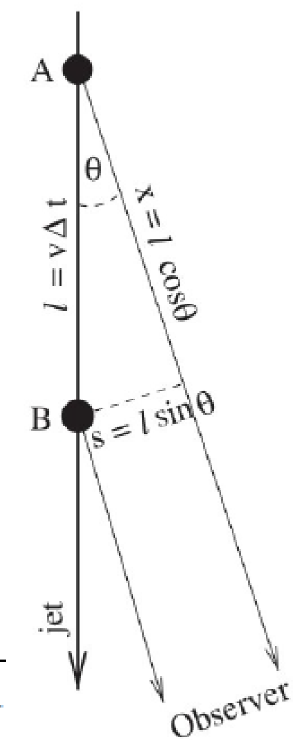
Physics of AGN jets

The relativistic Doppler factor: $\delta = \dfrac{1}{\Gamma(1-\beta_\Gamma \mu^{obs})}$ → Approaching sources are **boosted**, receding ones are **dimmed**.

Relativistic boosting: $F^{obs}_{\nu^{obs}} = F^{em}_{\nu^{em}}\,\delta^3$ and, in the case of power-law spectra: $F^{obs}_{\nu^{obs}} = F^{em}_{\nu^{obs}}\,\delta^{3+\alpha}$

Super-luminal motions (predicted by Special Relativity for sources approaching close to the line of sight):

$$v_{\perp,app} = \frac{s}{\Delta t_{obs}} = \frac{v\sin\vartheta}{(1-\beta_\Gamma\cos\vartheta)}$$



The maximum apparent velocity is achieved under the condition:

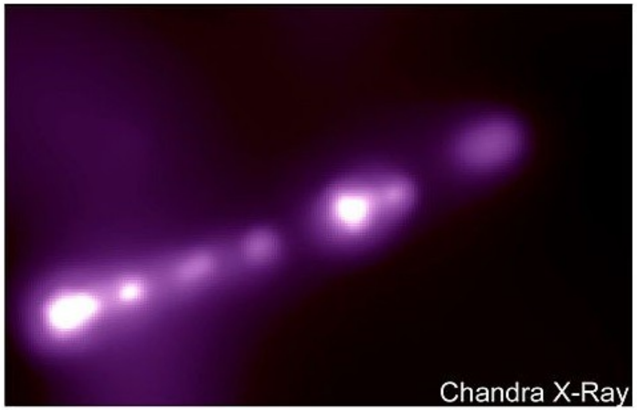
$$\frac{\mathrm{d}}{\mathrm{d}\cos\vartheta}\left(v_{\perp,app}\right) = 0 \quad \Longleftrightarrow \quad \cos\vartheta_{sl} = \beta_\Gamma$$

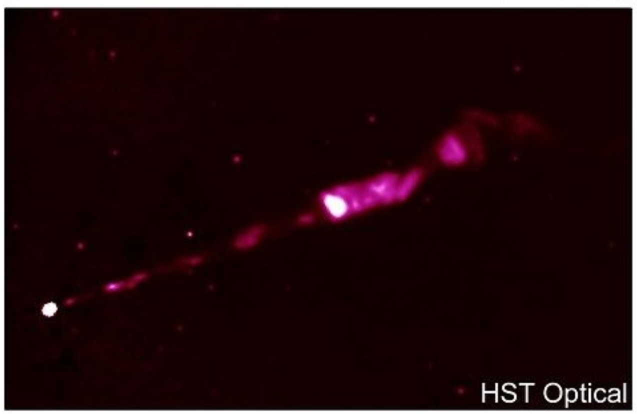
which defines a direction of maximal boosting:

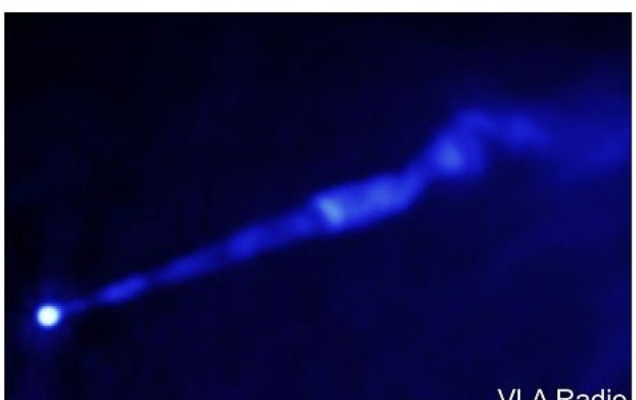
$$\sin\vartheta_{sl} = \sqrt{1-\cos^2\vartheta_{sl}} = \sqrt{1-\beta_\Gamma^2} = \frac{1}{\Gamma}$$

Finally it results that: $v_{\perp,app}(\vartheta_{sl}) = \dfrac{\frac{v}{\Gamma}}{1-\beta_\Gamma^2} = \Gamma v = c\beta_\Gamma\Gamma = c\sqrt{\Gamma^2-1}$

Physics of AGN jets







As it is predicted for relativistic charged particles in a **magnetic field *B***, we detect synchrotron spectrum radiation from the **radio frequencies** all the way up to **x-ray energies**.

The jets of nearby AGNs, like M87, can be resolved in their structure and they show a **knotty, multi-component** morphology.

The **high frequency** jet turns out to be composed by **distinct plasma blobs**, ejected by the nucleus at different epochs, rather than by a continuous stream. The **low frequency** jet appears **smoother**, with a rather complicated **spine-sheath structure** that reveals different kinematics.